

Møller's Energy-Momentum Complex for a Spacetime Geometry on a Noncommutative Curved D3-Brane

I. Radinschi · T. Grammenos

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Abstract Møller's energy-momentum complex is employed in order to determine the energy and momentum distributions for a spacetime described by a "generalized Schwarzschild" geometry in $(3 + 1)$ -dimensions on a noncommutative curved D3-brane in an effective, open bosonic string theory. The geometry considered is obtained by an effective theory of gravity coupled with a nonlinear electromagnetic field and depends only on the generalized (effective) mass and charge which incorporate corrections of first order in the noncommutativity parameter.

Keywords Møller's energy-momentum complex · Noncommutative string theory · Dp-branes

1 Introduction

The localization of energy and momentum in curved spacetime has been one of the most intricate issues and a subject of extensive research since the outset of general relativity. There has been a large number of attempts aiming at the localization of energy-momentum in this context, and they all have resulted in various (pseudotensorial) energy-momentum complexes. In order to solve the problem of energy and momentum localization, Einstein [1] was the first who proposed such a prescription followed, mainly, by Tolman [2], Papapetrou [3], Landau and Lifshitz [4], Bergmann and Thomson [5], Goldberg [6], Møller [7] and Weinberg [8]. With the exception of Møller's prescription, who proposed a new expression for an energy-momentum complex that can be utilized in any coordinate system, all the aforementioned energy-momentum complexes are restricted to quasi-Cartesian coordinates. However,

I. Radinschi
Department of Physics, "Gh. Asachi" Technical University, Iasi 700050, Romania
e-mail: radinschi@yahoo.com

T. Grammenos (✉)
Department of Mechanical and Industrial Engineering, University of Thessaly, 383 34 Volos, Greece
e-mail: thgramme@uth.gr

one should point out that, from the very beginning, the use of energy-momentum complexes has been heavily criticized on the grounds of their nontensorial character which could lead to rather dubious physical interpretations [9]. Furthermore, different localization prescriptions could lead to different energy distributions for the same spacetime geometry [10, 11], while due to their local character it was widely accepted that the total energy-momentum of the gravitational field could not be localized [12]. In 1973, Misner, Thorne and Wheeler [13] sustained that “anybody who looks for a magic formula for local gravitational energy-momentum is looking for the right answer to the wrong question”. This statement gives the actual meaning of the nonuniqueness of the pseudotensor for the energy-momentum. However, they concluded that the energy is indeed localizable only for spherical systems. In 1978, Cooperstock and Sarracino [14] showed that if the energy is localizable in spherical systems, then it is also localizable in any spacetime. In 1990, Bondi [15] maintained that “a non-localizable form of energy is not admissible in general relativity, because any form of energy contributes to gravitation and so its location can in principle be found”. Further, worth mentioning attempts to deal with the problems stemming from the concept of energy-momentum complex rely on the quasi-local energy approach [16–21], while there has been also a considerable amount of investigations on the concept of superenergy and the Bel-Robinson tensor [22–35]. The issue of the application of the pseudotensorial energy-momentum complexes for the localization of energy was almost abandoned until 1990, when Virbhadra and other scientists revived and enlightened the subject [36–46]. Since then a growing number of studies where different energy-momentum complexes are applied to a plethora of spacetime geometries has been published [47–76]. In 1996, Aguirregabiria, Chamorro and Virbhadra [77] established an important result showing that the Einstein, Tolman, Landau-Lifshitz, Papapetrou and Weinberg (henceforth ETLLPW) energy-momentum complexes essentially coincide if calculations are carried out in Kerr-Schild Cartesian coordinates for a general metric of the Kerr-Schild class. Furthermore, the results of the ETLLPW energy-momentum complexes were found to be identical with the older results of Penrose [78] and Tod [79], who had applied the notion of quasi-local mass. In 1999, in an attempt to support and stress the importance and usefulness of the concept of energy-momentum complex for the localization of energy, Chang, Nester and Chen [80, 81] showed that there exists a direct relationship of every energy-momentum complex with a Hamiltonian boundary term. Consequently, an energy-momentum complex provides a quasi-local, legitimate and acceptable expression. Work on this topic is still going on.

At this point it is worth noting several interesting results obtained in the last years in the study of the problem of energy localization in $2 + 1$ and $1 + 1$ dimensions by using energy-momentum complexes [82–89]. Further, a great deal of work has been done in the context of the tele-parallel version of gravity, where it is demonstrated that the energy-momentum complexes yield the same results as their tele-parallel versions [90–113].

In recent years, many physically satisfactory results were obtained by applying Møller’s prescription [114–152]. In this paper, we wish to emphasize the importance of the localization of energy using the same prescription, i.e., Møller’s energy-momentum complex. To this purpose, we extend our earlier study [153] and apply Møller’s prescription to evaluate the energy and momentum distributions for a metric describing a generalized Schwarzschild geometry in $(3 + 1)$ dimensions on a noncommutative D3-brane in the context of an effective, open bosonic string theory. The rest of the paper is organized as follows: In Sect. 2 the metric of the generalized Schwarzschild geometry and the general lines of its derivation are presented. Section 3 contains the definition of Møller’s energy-momentum complex, while in Sect. 4 we give the details of the calculations for the energy and momentum distributions.

Finally, Sect. 5 is devoted to a summary of the results obtained and to conclusions. The calculations have been performed using *Mathematica* and *Maple*, the latter having attached the GrTensor platform.

Conventions: throughout this paper we have used geometrized units ($c = 1, G = 1$), the metric’s Lorentzian signature is $(-, +, +, +)$, and Greek (Latin) indices are running from 0 to 3 (1 to 3).

2 The Generalized Schwarzschild Geometry

The nonperturbative approach to string theory has led to certain higher dimensional, extended objects called Dp-branes and defined as spacetime hypersurfaces onto which open strings can be attached and Dirichlet boundary conditions can be formulated [154]. At low energy scales or for slowly varying fields, the dynamics of Dp-branes is controlled by the effective Dirac-Born-Infeld (DBI) action, a natural generalization of the Dirac-Nambu-Goto action. The DBI action, obtained by one of the simplest nonpolynomial Lagrangeans invariant under reparametrizations of the world-volume, is based on a U(1) gauge field “living” on the world-volume swept out by the Dp-brane coupled to the inherent geometry of the world-volume.

In order to address the Hawking radiation phenomenon, the information loss paradox and, if possible, to formulate an effective theory of quantum gravity at the Planck scale, Kar and Majumdar [155] utilized the DBI-framework of a brane world-volume incorporating Einstein’s theory of general relativity coupled to Born-Infeld nonlinear electrodynamics. Then, considering the evolution of gravity on a noncommutative, curved D3-brane in an effective, open bosonic string theory setting (by using the Seiberg-Witten map [156], the authors transformed the ordinary gauge dynamics to a noncommutative U(1) gauge dynamics on the D3-brane) they obtained a generalized Schwarzschild geometry at the Planck scale. Essentially, in going towards the Planck regime, the authors combined an effective theory of gravity with a U(1) noncommutative gauge field and managed to proceed from a Schwarzschild-like geometry (described by (11) in our previous paper [153]), obtained with an ordinary spacetime in the classical regime, to a generalized Schwarzschild geometry on a noncommutative D3-brane, whereby a static gauge condition was used in order to incorporate a nontrivial induced metric on the brane world-volume. From a topological point of view, the generalized Schwarzschild geometry obtained describes a Euclidean manifold $\mathbb{R}^2 \times S^2$.

Thus, following Kar and Majumdar [155], we consider a Euclidean world-volume spanned by (x_1, x_2, x_3, x_4) with a signature $(+, +, +, +)$, where the Lorentzian signature may be obtained by an analytic continuation $x_4 \rightarrow it$, and the noncommutative D3-brane dynamics is obtained from the, appropriately modified, effective DBI-action given by

$$\hat{S}_{\text{DBI}} = T_D^{nc} \int d^4x \left(\sqrt{G} - \sqrt{G + 2\pi\alpha' \hat{F}} \right), \tag{1}$$

where T_D^{nc} is the noncommutative D3-brane tension, $G \equiv \det G_{\mu\nu}$, where $G_{\mu\nu} = g_{\mu\nu} - (Bg^{-1}B)_{\mu\nu}$ is the effective metric on the brane, $B_{\mu\nu}$ is a constant 2-form induced on the world-volume of the brane, and the noncommutative U(1) field strength on the world-volume is given by

$$\hat{F}_{\mu\nu} = \partial_\mu \hat{A}_\nu - \partial_\nu \hat{A}_\mu + \Theta^{\rho\lambda} \partial_\rho \hat{A}_\mu(x) \partial_\lambda \hat{A}_\nu(y) + O(\Theta^2), \tag{2}$$

where \hat{A}_μ is the noncommutative gauge potential and Θ is the noncommutativity parameter expressing the noncommutative corrections in the theory. Now, the action for a curved D3-brane becomes

$$\hat{S} = \frac{1}{16\pi} \int d^4x \sqrt{G} R + \hat{S}_{\text{DBI}} \tag{3}$$

with R the scalar curvature, while the appearing (string) gauge field corrections described by the higher order terms can vanish identically in a special combination of the gauge field. Hence, the effective noncommutative string action is finally obtained as

$$\hat{S} = \int d^4x \sqrt{G} \left(\frac{1}{16\pi} R - \frac{1}{4} G^{\mu\lambda} G^{\nu\rho} \hat{F}_{\mu\nu} * \hat{F}_{\lambda\rho} \right), \tag{4}$$

where the star $*$ denotes the Moyal product.

With a specific gauge choice for the effective metric $G_{\mu\nu}$, the Schwarzschild-like geometry ((11) in [153]) can be generalized to the effective noncommutative theory (4) at the Planck scale. In other words, a generalized semi-classical solution on a noncommutative curved D3-brane with the action (3) can be constructed using the effective metric $G_{\mu\nu}$. In the presence of a nonzero Neveu-Schwarz B-field (being equivalent to a constant magnetic field on the brane), the effective metric for the aforementioned geometry is finally given in spherical-polar coordinates by the following generalized Schwarzschild line element [155]:

$$ds^2 = - \left(1 - \frac{2M_{\text{eff}}}{r} - \frac{r_c^4}{r^4} + \frac{2M_{\text{eff}} r_c^4}{r^5} \right) dt^2 + \left(1 - \frac{2M_{\text{eff}}}{r} - \frac{r_c^4}{r^4} + \frac{2M_{\text{eff}} r_c^4}{r^5} \right)^{-1} dr^2 + \left(1 - \frac{r_c^4}{r^4} \right) r^2 d\theta^2 + \left(1 - \frac{r_c^4}{r^4} \right)^{-1} r^2 \sin^2 \theta d\phi^2 \tag{5}$$

with the effective parameters of the noncommutative formalism [157–159]

$$M_{\text{eff}} = M \left(1 - \frac{\Theta}{r^2} + O(\Theta^2) + \dots \right), \tag{6}$$

$$r_c^2 = Q \left(1 - \frac{\Theta}{2r^2} + O(\Theta^2) + \dots \right), \tag{7}$$

where M_{eff} is the generalized ADM mass, r_c^2 acts as an effective charge Q_{eff} , and Q is the electric charge. Keeping only first order terms in the noncommutativity parameter Θ , the line element (5) becomes:

$$ds^2 = - \left[1 - \frac{2M}{r} \left(1 - \frac{\Theta}{r^2} \right) - \frac{Q^2}{r^4} \left(1 - \frac{\Theta}{r^2} \right) + \frac{2M Q^2}{r^5} \left(1 - \frac{2\Theta}{r^2} \right) \right] dt^2 + \left[1 - \frac{2M}{r} \left(1 - \frac{\Theta}{r^2} \right) - \frac{Q^2}{r^4} \left(1 - \frac{\Theta}{r^2} \right) + \frac{2M Q^2}{r^5} \left(1 - \frac{2\Theta}{r^2} \right) \right]^{-1} dr^2 + \left[1 - \frac{Q^2}{r^4} \left(1 - \frac{\Theta}{r^2} \right) \right] r^2 d\theta^2 + \left[1 - \frac{Q^2}{r^4} \left(1 - \frac{\Theta}{r^2} \right) \right]^{-1} r^2 \sin^2 \theta d\phi^2. \tag{8}$$

As it is evident, the geometry is characterized only by the effective mass and the effective charge parameters. It is to be noted that the effective four-dimensional spacetime described

by the line element (5) may be obtained from a five-dimensional gravity with a small scale along the fifth dimension [157–159]. The effective theory solution described by the line element (5) breaks the spherical symmetry, it is of Petrov type I and describes a G_2 space (two Killing vectors) which is not maximally symmetric. Further, it has an obvious curvature singularity at $r = 0$ and two coordinate singularities at $r = r_c$ and $r = 2M_{\text{eff}}$. In fact, the noncommutative corrections may be viewed as a stretch in the event horizon of the object having this unusual exterior geometry.

In the classical regime, the corrections in the effective metric are small and may be ignored. Thus, in this case, the generalized geometry described by (8) is reduced to the classical Schwarzschild spacetime geometry.

3 Møller’s Energy-Momentum Complex

Arguing that the singling out of quasi-Cartesian coordinates in the Einstein energy-momentum complex is not satisfactory from the viewpoint of general relativity, Møller [7] searched for an expression for energy and momentum that would be independent from any particular coordinate system. To that purpose, Møller sought for a quantity with identically vanishing divergence that could be added to Einstein’s energy-momentum complex so that their sum could be transformed as a tensor for spatial transformations. Møller’s quantity should be form-invariant and dependent on the metric tensor as well as on its first and second derivative. Further, in linear transformations it should behave as a tensor density satisfying certain conditions (see, e.g., [160] for details). Employing this procedure, Møller arrived at the following energy-momentum complex in a four-dimensional background:

$$J_v^\mu = \frac{1}{8\pi} \xi_{v,\lambda}^{\mu\lambda}, \tag{9}$$

where Møller’s superpotential $\xi_v^{\mu\lambda}$ has the form

$$\xi_v^{\mu\lambda} = \sqrt{-g} \left(\frac{\partial g_{v\sigma}}{\partial x^\kappa} - \frac{\partial g_{v\kappa}}{\partial x^\sigma} \right) g^{\mu\kappa} g^{\lambda\sigma} \tag{10}$$

with the antisymmetry property

$$\xi_v^{\mu\lambda} = -\xi_v^{\lambda\mu}. \tag{11}$$

It is easily seen that Møller’s energy-momentum complex satisfies the local conservation equation

$$\frac{\partial J_v^\mu}{\partial x^\mu} = 0, \tag{12}$$

where J_0^0 gives the energy density and J_i^0 gives the momentum density components.

Finally, in Møller’s prescription the energy and momentum in a four-dimensional background are obtained by

$$P_v = \iiint J_v^0 dx^1 dx^2 dx^3. \tag{13}$$

Specifically, the energy of a physical system is given by

$$E = \iiint J_0^0 dx^1 dx^2 dx^3. \tag{14}$$

As it was pointed out in the beginning of this section, the calculations are not anymore restricted to quasi-Cartesian coordinates, but can be utilized in any coordinate system. Indeed, Lessner [161] has concluded that Møller’s energy-momentum complex is a powerful expression for the localization of energy and momentum in general relativity.

4 Energy and Momentum Density Distributions

For the metric under consideration given by the line element (8), we obtain eight non-vanishing Møller superpotentials, with the antisymmetric property (11) taken into account:

$$\xi_2^{21} = -\xi_2^{12} = -\frac{2 \sin \theta [r^6 + Q^2(r^2 - 2\Theta)]}{r^6[r^6 + Q^2(-r^2 + \Theta)]} \{r^7 + Q^2(-r^3 + r\Theta) + 2M[-r^6 + Q^2(r^2 - 2\Theta) + r^4\Theta]\}, \tag{15}$$

$$\xi_3^{31} = -\xi_3^{13} = -\frac{2 \sin \theta [r^6 + Q^2(-3r^2 + 4\Theta)]}{r^6[r^6 + Q^2(-r^2 + \Theta)]} \{r^7 + Q^2(-r^3 + r\Theta) + 2M[-r^6 + Q^2(r^2 - 2\Theta) + r^4\Theta]\}, \tag{16}$$

$$\xi_3^{32} = -\xi_3^{23} = \frac{2r^6}{-r^6 + Q^2r^2 - \Theta Q^2} \cos \theta, \tag{17}$$

$$\xi_0^{01} = -\xi_0^{10} = \frac{-2 \sin \theta}{r^6} [r^2(2Q^2r + Mr^4 - 5MQ^2) + \Theta(-3Q^2r - 3Mr^4 + 14MQ^2)]. \tag{18}$$

By substituting the superpotentials given by (15–18) into (9) we obtain the energy density distribution

$$J_0^0 = \frac{\sin \theta}{4\pi r^7} [\Theta(-15Q^2r - 6Mr^4 + 84MQ^2) + 2Q^2r^2(3r - 10M)], \tag{19}$$

while all the momentum density distributions vanish:

$$J_i^0 = 0. \tag{20}$$

Thus, all the momentum components are zero. Inserting (19) into (14) and evaluating the integral, we obtain the energy contained in a “sphere” of radius R :

$$E = M - \frac{1}{R^6} \{R^2 Q^2 (2R - 5M) - \Theta [3R^4 M + Q^2 (3R - 14M)]\}. \tag{21}$$

Evidently, the energy contained in this “sphere” depends only on the radius R , the mass M and the charge Q generalized by the noncommutativity parameter Θ and given by equations (6, 7). In the zero-charge limit $Q \rightarrow 0$, one obtains from (21):

$$E = M \left(1 + \frac{3\Theta}{R^2} \right), \tag{22}$$

while, in the case there are not any noncommutativity corrections, i.e. $\Theta = 0$, (21) yields

$$E = M - \frac{2Q^2}{R^3} \left(1 - \frac{5}{2} \frac{M}{R} \right). \tag{23}$$

Taking the asymptotic limit $R \rightarrow \infty$, one derives from (21) that the energy equals the (ADM) mass M :

$$E = M, \quad (24)$$

a result that can also be obtained from (21) when Q and Θ vanish simultaneously. However, one must point out that the energy E equals the (ADM) mass M also when $\Theta = 0$ and $R = \frac{5}{2}M$. No reasonable explanation could be found for this result that was obtained also in [153]. An explanation of this rather unexpected result, namely that the energy E equals the (ADM) mass M for a finite radius R , could be possibly offered by a detailed study of the properties of the spacetime geometry described by the line element (5).

5 Conclusions

The different attempts made to develop a generally accepted expression for the energy-momentum density yield a plethora of energy-momentum definitions. Indeed, for solving the problem of energy and momentum localization many attempts have been made in the past, but it still remains an important open issue to be settled. Even though the energy-momentum complexes are apparently useful for the localization of energy, there are still doubts that these prescriptions could give acceptable results for any given space-time for reasons discussed in Sect. 1. This is, among other things, due to the fact that many of the energy-momentum complexes (e.g., ELLTPW and Bergmann-Thomson) are restricted to the use of particular coordinate systems and only the Møller energy-momentum complex allows one to calculate the energy distribution in any coordinate system. Nevertheless, the results obtained by several authors [47–153] have demonstrated that the energy-momentum complexes are useful concepts for evaluating the energy and momentum distributions in general relativity.

In this work, the energy and momentum distributions for a semiclassical generalized Schwarzschild solution in (3 + 1) dimensions that reduces to the classical Schwarzschild solution in the classical regime and is derived by an effective, open bosonic string theory on a noncommutative curved D3-brane in a static gauge, with a mass and charge being noncommutatively corrected to first order in the noncommutativity parameter Θ (giving a generalized mass and a generalized charge), have been calculated by using Møller's energy-momentum complex. There are totally eight nonvanishing Møller superpotentials and all momentum components are found to be zero. The total energy contained in a "sphere" is computed, while its value is obtained in the zero-charge limit, in the asymptotic limit, and when there are no noncommutativity corrections.

Finally, an investigation of the considered four-dimensional exterior geometry described by the line element (5) showed that it is of Petrov type I. However, one must point out that this classification might change by some proper combination of mass and charge. Further, the metric describes a G2 space (with $\frac{\partial}{\partial t}$ and $\frac{\partial}{\partial \phi}$ as Killing vectors) which is not maximally symmetric (non-zero Weyl projective tensor) and it is not an Einstein space. A calculation and examination of the curvature invariants for this metric, even though computationally tedious, would be of interest for a clarification of further properties this spacetime geometry might have.

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References

1. Einstein, A.: *Preuss. Akad. Wiss. Berlin* **47**, 778 (1915). Addendum: *Preuss. Akad. Wiss. Berlin* **47**, 799 (1915)
2. Tolman, R.C.: In: *Relativity, Thermodynamics and Cosmology*, p. 227. Oxford University Press, London (1934)
3. Papapetrou, A.: *Proc. R. Ir. Acad. A* **52**, 11 (1948)
4. Landau, L.D., Lifshitz, E.M.: In: *The Classical Theory of Fields*, p. 317. Addison-Wesley, Reading (1951)
5. Bergmann, P.G., Thomson, R.: *Phys. Rev.* **89**, 400 (1953)
6. Goldberg, J.N.: *Phys. Rev.* **111**, 315 (1958)
7. Møller, C.: *Ann. Phys. (NY)* **4**, 347 (1958)
8. Weinberg, S.: In: *Gravitation and Cosmology: Principles and Applications of General Theory of Relativity*, p. 165. Wiley, New York (1972)
9. Chandrasekhar, S., Ferrari, V.: *Proc. R. Soc. Lond. A* **435**, 645 (1991)
10. Bergqvist, G.: *Class. Quantum Gravity* **9**, 1753 (1992)
11. Bergqvist, G.: *Class. Quantum Gravity* **9**, 1917 (1992)
12. Chen, C.M., Nester, J.M.: *Class. Quantum Gravity* **16**, 1279 (1999)
13. Misner, C.W., Thorne, K.S., Wheeler, J.A.: In: *Gravitation*, p. 467. Freeman, New York (1973). See also p. 603
14. Cooperstock, F.I., Sarracino, R.S.: *J. Phys. A: Math. Gen.* **11**, 877 (1978)
15. Bondi, H.: *Proc. R. Soc. Lond. A* **427**, 249 (1990)
16. Brown, J.D., York, J.W.: *Phys. Rev. D* **47**, 1407 (1993)
17. Hayward, S.A.: *Phys. Rev. D* **49**, 831 (1994)
18. Komar, A.: *Phys. Rev.* **113**, 934 (1959)
19. Brown, J.D., Lau, S.R., York, J.W.: *Phys. Rev. D* **55**, 1977 (1997)
20. Yau, S.-T.: *Adv. Theor. Math. Phys.* **5**, 755 (2001)
21. Liu, C.-C.M., Yau, S.-T.: *Phys. Rev. Lett.* **90**, 231102 (2003)
22. Senovilla, J.M.M.: *Class. Quantum Gravity* **17**, 2799 (2000)
23. Bel, L.: *Comm. Math. Phys.* **138**, 59 (1962)
24. Robinson, I.: *Class. Quantum Gravity* **14**, A331 (1997)
25. Bonilla, M.A.G., Senovilla, J.M.M.: *Phys. Rev. Lett.* **11**, 783 (1997)
26. Bonilla, M.A.G., Senovilla, J.M.M.: *Gen. Relativ. Gravit.* **29**, 91 (1997)
27. Senovilla, J.M.M.: In: Martin, J., Ruiz, E., Atrio, F., Molina, A. (eds.) *Gravitation and Relativity in General*. World Scientific (1999). gr-qc/9901019
28. Senovilla, J.M.M.: gr-qc/9912050 (1999)
29. Teyssandier, P.: gr-qc/9905080 (1999)
30. Garecki, J.: *Ann. Phys.* **10**, 911 (2001)
31. Dabrowski, M.P., Garecki, J.: *Class. Quantum Gravity* **19**, 1 (2002)
32. Lazkog, R., Senovilla, J.M.M., Vera, R.: *Class. Quantum Gravity* **20**, 4135 (2003)
33. Garecki, J.: *Class. Quantum Gravity* **22**, 4051 (2005)
34. Senovilla, J.M.M.: *Class. Quantum Gravity* **23**, 7143 (2006)
35. Saha, B., Rikhvitsky, V.: *Mod. Phys. Lett. A* **21**, 847 (2006)
36. Virbhadra, K.S.: *Phys. Rev. D* **41**, 1086 (1990)
37. Virbhadra, K.S.: *Phys. Rev. D* **42**, 1066 (1990)
38. Virbhadra, K.S.: *Phys. Rev. D* **42**, 2919 (1990)
39. Virbhadra, K.S.: *Pramana J. Phys.* **38**, 31 (1992)
40. Virbhadra, K.S.: *Pramana J. Phys.* **45**, 215 (1995)
41. Virbhadra, K.S.: *Phys. Rev. D* **60**, 104041 (1999)
42. Rosen, N., Virbhadra, K.S.: *Gen. Relativ. Gravit.* **25**, 429 (1993)
43. Virbhadra, K.S., Parikh, J.C.: *Phys. Lett. B* **317**, 312 (1993)
44. Virbhadra, K.S., Parikh, J.C.: *Phys. Lett. B* **331**, 302 (1994). Erratum: *Phys. Lett. B* **340**, 265 (1994)
45. Chamorro, A., Virbhadra, K.S.: *Pramana J. Phys.* **45**, 181 (1995)
46. Chamorro, A., Virbhadra, K.S.: *Int. J. Mod. Phys. D* **5**, 251 (1997)
47. Garecki, J.: *Class. Quantum Gravity* **22**, 4051 (2005)
48. Aydogdu, O.: *Int. J. Mod. Phys. D* **15**, 459 (2006)
49. Sharif, M., Fatima, T.: *Nuovo Cimento B* **120**, 533 (2005)
50. Salti, M.: *Nuovo Cimento B* **120**, 53 (2005)
51. Gad, R.M.: *Gen. Relativ. Gravit.* **38**, 417 (2006)
52. Gad, R.M.: *Nuovo Cimento B* **121**, 161 (2006)
53. Radinschi, I., Ciobanu, B.: *Rom. J. Phys.* **50**, 1223 (2005)

54. Ciobanu, B., Radinschi, I.: Rom. J. Phys. **51**, 1003 (2006)
55. Radinschi, I., Ciobanu, B.: AIP Conf. Proc. **895**, 329 (2007). gr-qc/0608029
56. Radinschi, I.: Acta Phys. Slovaca **49**, 789 (1999)
57. Radinschi, I.: Acta Phys. Slovaca **50**, 609 (2000)
58. Radinschi, I.: Fizika B **9**, 43 (2000)
59. Radinschi, I.: Nuovo Cimento B **115**, 501 (2000)
60. Radinschi, I.: Chin. J. Phys. **39**, 231 (2001)
61. Yang, I.-C., Radinschi, I.: Chin. J. Phys. **41**, 326 (2003)
62. Yang, I.-C.: Chin. J. Phys. **38**, 1040 (2000)
63. Halpern, P.: gr-qc/0705.0720 (2007)
64. Halpern, P.: Mod. Phys. Lett. A **21**, 495 (2006)
65. Halpern, P.: Astrophys. Space Sci. **306**, 279 (2006)
66. Aydogdu, O.: Fortschr. Phys. **54**, 246 (2006)
67. Aydogdu, O., Salti, M.: Prog. Theor. Phys. **115**, 63 (2006)
68. Salti, M., Havare, A.: gr-qc/0502058 (2005)
69. Gad, R.M.: Astrophys. Space Sci. **295**, 451 (2005)
70. Gad, R.M.: Mod. Phys. Lett. A **19**, 1847 (2004)
71. Gad, R.M.: Gen. Relativ. Gravit. **38**, 417 (2006)
72. Patashnick, O.: Int. J. Mod. Phys. D **14**, 1607 (2005)
73. Xulu, S.S.: gr-qc/0702066 (2007)
74. Xulu, S.S.: Int. J. Mod. Phys. A **15**, 2979 (2000)
75. Xulu, S.S.: Int. J. Theor. Phys. **39**, 1153 (2000)
76. Xulu, S.S.: Int. J. Mod. Phys. A **15**, 4849 (2000)
77. Aguirregabiria, J.M., Chamorro, A., Virbhadra, K.S.: Gen. Relativ. Gravit. **28**, 1393 (1996)
78. Penrose, R.: Proc. R. Soc. Lond. A **381**, 53 (1982)
79. Tod, K.P.: Proc. R. Soc. Lond. A **388**, 457 (1983)
80. Chang, C.C., Nester, J.M., Chen, C.M.: Phys. Rev. Lett. **83**, 1897 (1999)
81. Chen, C.M., Liu, J.L., Nester, J.M.: gr-qc/0705.1080 (2007)
82. Vagenas, E.C.: Int. J. Mod. Phys. A **18**, 5949 (2003)
83. Vagenas, E.C.: Int. J. Mod. Phys. A **18**, 5781 (2003)
84. Vagenas, E.C.: J. High Energy Phys. **0307**, 046 (2003)
85. Yang, I.-C., Radinschi, I.: AIP Conf. Proc. **895**, 325 (2007). gr-qc/0309130
86. Vagenas, E.C.: Int. J. Mod. Phys. D **14**, 573 (2005)
87. Grammenos, T.: Mod. Phys. Lett. A **20**, 1741 (2005)
88. Radinschi, I., Grammenos, T.: Int. J. Mod. Phys. A **21**, 2853 (2006)
89. Radinschi, I., Yang, I.-C.: gr-qc/0702105 (2007)
90. Nashed, G.G.L.: Phys. Rev. D **66**, 064015 (2002)
91. Nashed, G.G.L.: Nuovo Cimento B **117**, 521 (2002)
92. Nashed, G.G.L.: Mod. Phys. Lett. A **22**, 1047 (2007)
93. Salti, M., Havare, A.: Int. J. Mod. Phys. A **20**, 2169 (2005)
94. Salti, M.: Mod. Phys. Lett. A **20**, 2175 (2005)
95. Salti, M.: Acta Phys. Slovaca **55**, 563 (2005)
96. Salti, M.: gr-qc/0607116 (2006)
97. Salti, M.: Astrophys. Space Sci.. gr-qc/0607011 (2006, to appear)
98. Salti, M.: Int. J. Mod. Phys. D **15**, 695 (2006)
99. Aydogdu, O., Salti, M.: Astrophys. Space Sci. **302**, 61 (2006)
100. Aydogdu, O., Salti, M.: Astrophys. Space Sci. **299**, 227 (2005)
101. Aydogdu, O., Salti, M.: gr-qc/0602012 (2006)
102. Aydogdu, O., Salti, M., Korunur, M.: Acta Phys. Slovaca **55**, 537 (2005)
103. Aydogdu, O.: Int. J. Mod. Phys. D **15**, 459 (2006)
104. Sharif, M., Amir, M.J.: Gen. Relativ. Gravit. **39**, 989 (2007)
105. Sharif, M., Amir, M.J.: Gen. Relativ. Gravit. **38**, 1735 (2006)
106. Sharif, M., Amir, M.J.: Mod. Phys. Lett. A **22**, 425 (2007)
107. Sharif, M., Amir, M.J.: gr-qc/0704099 (2007)
108. Sousa, A.A., Moura, J.S., Pereira, R.B.: gr-qc/0702109 (2007)
109. So, L.L., Vargas, T.: Chin. J. Phys. **43**, 901 (2005)
110. Korunur, M.: gr-qc/0608014 (2006)
111. Aygun, S., Aygun, M., Tarhan, I.: gr-qc/0607126 (2006)
112. Korunur, M., Salti, M., Aydogdu, O., Acikgoz, I.: gr-qc/0607117 (2006)
113. Aygun, S., Baysal, H., Tarhan, I.: gr-qc/0607109 (2006)
114. Aygun, S.: Acta Phys. Polonica B **38**, 73 (2007)

115. Aygun, S., et al.: *Acta Phys. Polonica B* **37**, 2781 (2006)
116. Salti, M., Aydogdu, O.: *Eur. Phys. J. C* **47**, 247 (2006)
117. Salti, M., Aydogdu, O.: *Int. J. Theor. Phys.* **45**, 1891 (2006)
118. Salti, M., Aydogdu, O.: *Int. J. Theor. Phys.* **45**, 2481 (2006)
119. Salti, M., Aydogdu, O.: *Found. Phys. Lett.* **19**, 269 (2006)
120. Salti, M.: *Czech. J. Phys.* **56**, 177 (2006)
121. Salti, M.: *Mod. Phys. Lett. A* **20**, 2175 (2005)
122. Salti, M., et al.: *J. High Energy Phys.* **0612**, 078 (2006)
123. Aydogdu, O.: *Int. J. Mod. Phys. A* **21**, 3845 (2006)
124. Aydogdu, O., Salti, M.: *Pramana J.* **67**, 239 (2006)
125. Aydogdu, O., Salti, M.: *Astrophys. Space Sci.* **299**, 227 (2005)
126. Aydogdu, O., Salti, M., et al.: *Found. Phys. Lett.* **19**, 709 (2006)
127. Nashed, G.G.L.: *Nuovo Cimento B* **119**, 967 (2005)
128. Gad, R.M.: *Astrophys. Space Sci.* **302**, 141 (2006)
129. Gad, R.M.: *Astrophys. Space Sci.* **295**, 459 (2005)
130. Gad, R.M.: *Astrophys. Space Sci.* **293**, 453 (2004)
131. Radinschi, I., Yang, I.-C.: *Fizika B* **14**, 311 (2005)
132. Radinschi, I.: In: Reimer, A. (ed.) *Quant. Cosmology Research Trends. Horizons in World Physics*, vol. 246. Nova Science Publ. (2005). gr-qc/0302047
133. Radinschi, I.: gr-qc/0203084 (2002)
134. Radinschi, I.: *Mod. Phys. Lett. A* **16**, 673 (2001)
135. Radinschi, I.: *Mod. Phys. Lett. A* **15**, 2171 (2000)
136. Yang, I.-C., Radinschi, I.: *Mod. Phys. Lett. A* **17**, 1159 (2002)
137. Yang, I.-C., Radinschi, I.: *Chin. J. Phys.* **42**, 40 (2004)
138. Vagenas, E.C.: *Mod. Phys. Lett. A* **21**, 1947 (2006)
139. Vagenas, E.C.: *Mod. Phys. Lett. A* **19**, 213 (2004)
140. Sharif, M., Fatima, T.: *Int. J. Mod. Phys. A* **20**, 4309 (2005)
141. Sharif, M., Fatima, T.: *Astrophys. Space Sci.* **302**, 217 (2006)
142. Sharif, M., Azam, M.: *Int. J. Mod. Phys. A* **22**, 1935 (2007)
143. Aygun, S., et al.: gr-qc/07040525 (2007)
144. Aygun, S., Baysal, H., Tarhan, I.: gr-qc/0608024 (2006)
145. So, L.L., Nester, J.M.: gr-qc/0612061 (2006)
146. Binbay, F., Pirinccioglu, N., Salti, M., Aydogdu, O.: gr-qc/0607095 (2006)
147. Yang, I.-C., Jeng, C.-C.: gr-qc/0605145 (2006)
148. Yang, I.-C., Tsai, C.H.: gr-qc/0503118 (2005)
149. Havare, A., Salti, M., Yetkin, T.: gr-qc/0502057 (2005)
150. Havare, A., Korunur, M., Salti, M.: *Astrophys. Space Sci.* **301**, 43 (2006)
151. Xulu, S.S.: *Astrophys. Space Sci.* **283**, 23 (2003)
152. Xulu, S.S.: *Mod. Phys. Lett. A* **15**, 1151 (2000)
153. Grammenos, T., Radinschi, I.: *Int. J. Theor. Phys.* **46**, 1055 (2007)
154. Johnson, C.V.: *D-Branes*. Cambridge University Press, Cambridge (2003)
155. Kar, S., Majumdar, S.: *Int. J. Mod. Phys. A* **21**, 6087 (2006)
156. Seiberg, N., Witten, E.: *J. High Energy Phys.* **9909**, 032 (1999)
157. Kar, S.: *J. High Energy Phys.* **10**, 052 (2006)
158. Kar, S.: *Phys. Rev. D* **74**, 126002 (2006)
159. Kar, S.: *Phys. Rev. D* **74**, 066003 (2006)
160. Xulu, S.S.: Ph.D. thesis. hep-th/0308070 (2003)
161. Lessner, G.: *Gen. Relativ. Gravit.* **28**, 527 (1996)